NEW LIMITS ON THE POLARIZED ANISOTROPY OF THE COSMIC MICROWAVE BACKGROUND AT SUBDEGREE ANGULAR SCALES

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ABSTRACT

We update the limit from the 90 GHz Princeton *IQU* Experiment ground-based polarimeter on the magnitude of any polarized anisotropy of the cosmic microwave radiation. With a second year of data, we have now limited both Q and U on a ring of 1° radius. The window functions are broad: for *E*-mode polarization, the effective l is $\langle l_E \rangle = 191^{+143}_{-132}$. We find that the *E*-mode signal can be no greater than 8.4 μK (95% CL), assuming no *B*-mode polarization. Limits on a possible *B*-mode signal are also presented.

Subject headings: cosmic microwave background — cosmology: observations — early universe

1. INTRODUCTION

The current limits on cosmic microwave background (CMB) polarization restrict the amplitude of its fluctuations to less than 10 μK at 95% confidence level (CL). At large angular scales, Keating et al. (2001) limit the amplitude to 8 μK. At subdegree angular scales, the constraint from Hedman et al. (2001, hereafter H01) is 10 μK, while at arcminute scales, Subrahmanyan et al. (2000) set a limit of 10 μK. Present estimates of the peak polarized fluctuation amplitude are ~6 μK at an angular scale of ~0°.2 ($l \sim 950$). These estimates are based on parameters gleaned from CMB temperature anisotropy measurements (e.g., Pryke et al. 2002; Jaffe et al. 2001; Wang, Tegmark, & Zaldarriaga 2002). While CMB polarization has yet to be detected, its characterization will complement CMB temperature anisotropy data and impact our understanding of gravitational waves from the inflationary epoch (e.g., Turner 1997; Caldwell, Kamionkowski, & Wadley 1999), peculiar velocities at the surface of last scattering (Zaldarriaga & Harari 1995), the nature of primordial perturbations (e.g., Spergel & Zaldarriaga 1997), primordial magnetic fields (Kosowsky & Loeb 1996), and cosmological parity violation (Lue, Wang, & Kamionkowski 1999). Here we report improved limits derived from new data from the 2001 observing season of the Princeton *IQU* Experiment (PIQUE) at 90 GHz. We combine the new data with data from the first observing season and also present a reanalysis of those earlier data. These data pass extensive checks for systematic contamination. Future publications will report results from a 40 GHz polarimeter also deployed during the 2001 observing season and give details of the instrument.

2. INSTRUMENT, OBSERVATIONS, AND CALIBRATION

PIQUE has been described previously (H01). The results reported here are from PIQUE's broadband 90 GHz correlation polarimeter, which underilluminates a 1.2 m off-axis parabola (Wollack et al. 1997), resulting in a beam size of 0°.235. The 84–100 GHz bandpass is divided into three subbands called S0, S1, and S2 (H01). Observations are made of a ring of

radius 1° around the north celestial pole; the telescope site is Princeton, New Jersey.

The polarimeter observed the sky from 2000 January 19 to 2000 April 2 and from 2000 December 19 to 2001 February 28. These two observing seasons yielded 810 and 660 hr of raw data, respectively.

The scanning strategy is designed to permit null tests for checking sensitivity to systematics. During both observing seasons, the telescope alternated between two azimuth positions at fixed elevation. Data from the two positions are differenced to remove sensitivity to constant offsets. For the first season, these azimuth positions were ±0°.93 and the elevation was 41°.0. The telescope therefore measured $\mp Q$ (as defined by the IAU) for two regions separated by 6 hr in right ascension on the ring of declination 89°. For the second season, the azimuth positions were ±1°.31, with elevation 40°.3, so the telescope measured $+U$ for two regions separated by 12 hr in right ascension on the same ring. The azimuth chop period was 13 s until 2001 January 23, at which point it was doubled.

The polarimetry channels are calibrated to 10% using a nutating aluminum flat (H01; Staggs et al. 2002). Constant elevation scans of Jupiter are used to determine pointing accuracy and map the beams. For the second observing season, the measured beam FWHMs are 0°.235(7) in coelevation and 0°.233(7) in elevation, in agreement with measurements from the first season. The absolute pointing offsets in elevation are smaller than 0°.03. However, early in the second observing season, the encoder suffered a misalignment so that the azimuth offset increased from +0°.03 to +0°.08. (No concomitant change in beam shape was observed.) Note that we are able to neglect this in the analysis because the overlap between the ideal and misaligned beams is still 85%. In fact, simulations indicate the misalignment has less than a 2% effect on our derived limit.

3. DATA REDUCTION

The 1470 hr of data from the two observing seasons include 383 hr of data taken while the telescope was slewing between the desired scan positions, 339 hr of data corrupted by known electromechanical failures, and 59 hr of data in isolated fragments less than 12 hr in length. The remaining 685 hr of data contain 186 hr of data corrupted by meteorological phenomena (clouds), which are identified using the selection criteria described below.

As discussed in H01, the correlation polarimeter suffers a small sensitivity ($\lesssim 0.5\%$) to total power because signals reflected from the input of the amplifier in one arm can couple

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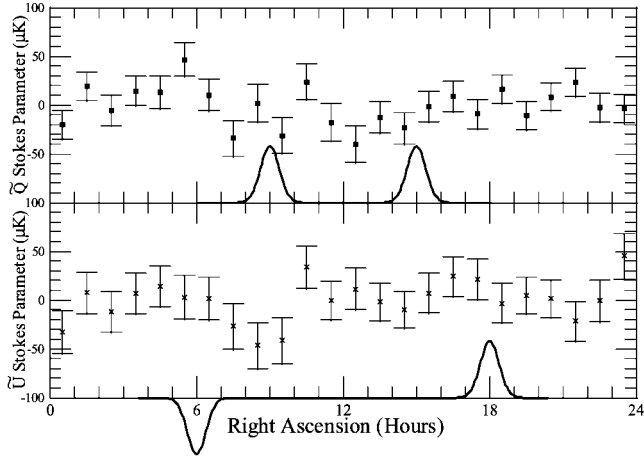
FIG. 1.—Binned data, in thermodynamic units, for each observation season. The three frequency channels have been co-added, and the 1 σ errors include small correlations among the channels. For clarity, only 24 approximately beam-sized bins are shown: the actual analysis uses 144 independent bins. Owing to PIQUE's differencing strategies, a 1 Jy point source with 20% polarization located at $\alpha = 12^h$, $\delta = 89°$ generates the point-source responses shown in the two panels. No such point source is known, and Toffolatti et al. (1998) estimate there are fewer than 200 point sources brighter than 1 Jy in the entire sky at 90 GHz.

into the other arm through the orthomode transducer. Therefore, the distributions of correlation coefficients

$$C_N = \frac{\langle S_i S_j \rangle}{\langle S_i^2 \rangle \langle S_j^2 \rangle} \qquad (1)$$

between pairs of polarimetry channels (i, j) display large positive tails due to periods of rapid atmospheric fluctuations. Data corrupted by clouds are removed by requiring the coefficients to be less than certain thresholds. Such selection criteria are determined based on a data set designed to be insensitive to real astronomical polarized signals: the quadrature data (for which data from each scan position are split into halves and differenced). The C_N are generated as averages over N chops; varying N varies the timescale probed. For purely Gaussian noise, the shapes of the resulting distributions of C_N for the whole data set depend only on N. In order to avoid using additional cutting measures (such as the 6 hr null test used in H01), the C_N-selection technique has been refined from that used for H01. Here two timescales are used rather than one. First we calculate the average correlation coefficients for segments of the time series 40–70 minutes long (specifically, $N = 200$). Segments with coefficients larger than a threshold of 0.20 are removed. Next, coefficients are calculated for $N = 50$, for which the cut threshold is 0.32. These thresholds are at 2.5 σ and 2 σ and are chosen so that either cut alone removes ∼20% of the data. The combined cuts remove 27% of the data. Null test results are not sensitive to the exact values for the thresholds.

The 307 (192) hr of data surviving the cuts detailed above from the first (second) observing season are parsed into 144 bins based on the local sidereal time (LST) when the data were taken, following the same procedures outlined in H01. These data are plotted in Figure 1. Offsets on the order of a few hundred microkelvins are removed from each polarimetry channel for each "deployment" (a period of >12 hr bracketed by periods when the instrument was tarped). The results are not sensitive to the exact number of offsets removed. Table 1

TABLE 1
RESULTS OF χ^2 CONSISTENCY TESTS

Year	Test	S0[a]	S1[a]	S2[a]
2000	Quadrature[b]	0.96	0.25	0.19
	H1 − H2[c]	0.18	0.14	0.33
	Pattern[d]	0.46	0.26	0.14
	Si − Sj[e]	0.62	0.83	0.83
2001	Quadrature[b]	0.38	0.61	0.28
	H1 − H2[c]	0.17	0.61	0.58
	Pattern[d]	0.80	0.12	0.38
	Si − Sj[e]	0.47	0.66	0.04

[a] Each numerical entry gives the probability of exceeding the χ^2 for the given frequency channel.

[b] The quadrature test uses data from each scan position (east and west) split into two halves and differenced to yield the quantity $(E_1 − E_2)/2 + (W_1 − W_2)/2$.

[c] Data from the second half of each season are subtracted from the first half.

[d] Pattern nulls are generalizations of the 6 hr null test from H01 and are data sets constructed from the various combinations of the data that should be zero given the differencing scan strategy. If d_t is the measured signal at LST t in hours, then for 2000 these combinations are $d_t − d_{t+6} + d_{t+12} − d_{t+18}$, while for 2001 these combinations are $d_t + d_{t+12}$.

[e] Data differenced between two channels. The column entries are S0 − S1, S0 − S2, and S1 − S2.

presents the results of the null tests described in H01, using the new selection criteria. The χ^2 distribution of these null tests is consistent with noise, demonstrating that the data do not suffer from residual atmospheric contamination.

4. DATA ANALYSIS

The likelihood of a model given a data vector x is $L \propto \exp(−x^T C^{-1} x/2)/|C|^{1/2}$, where the covariance matrix C sums both theoretical correlations from the model (signal) and correlations specific to the experiment (noise): $C = S + N$. The analysis of the combined 2000–2001 data uses x with 864 elements comprising 144 spatial pixels measured in three different frequency bands over the course of two observing seasons. The noise matrix N encodes the variances for each pixel for each channel and also accounts for interchannel correlations from both atmospheric fluctuations weakly coupled into the polarimeter channels and correlated gain fluctuations in the cryogenic amplifiers. The interchannel correlation coefficients are ≤8% on average and smaller for the two most sensitive channels. (S2 has just 20% of the total weight.) The noise matrix does not include pixel-pixel correlations, since no such correlations are observed in the time series data. An independent analysis using a 288-element data vector, for which data from the three frequency channels are combined for each spatial pixel, with errors calculated to account for the interchannel correlations, yields consistent results.

Since PIQUE measured Q in 2000 and U in 2001, the signal covariance matrix takes the form

$$S = \begin{pmatrix} \langle \tilde{Q}\tilde{Q} \rangle & \langle \tilde{Q}\tilde{U} \rangle \\ \langle \tilde{Q}\tilde{U} \rangle & \langle \tilde{U}\tilde{U} \rangle \end{pmatrix}, \qquad (2)$$

where $\langle \tilde{Q}_i \tilde{Q}_j \rangle$ (or $\langle \tilde{U}_i \tilde{U}_j \rangle$) represents the theoretical correlation between two spatial pixels from the 2000 (or 2001) data set and $\langle \tilde{Q}_i \tilde{U}_j \rangle$ encodes correlations between pixels from different years. Note that the $\tilde{Q}$ are sums of Q separated by 6 hr in right ascension, and the $\tilde{U}$ are differences of U separated by 12 hr in right ascension. The expression for $\langle \tilde{Q}_i \tilde{Q}_j \rangle$ is given in H01. Following
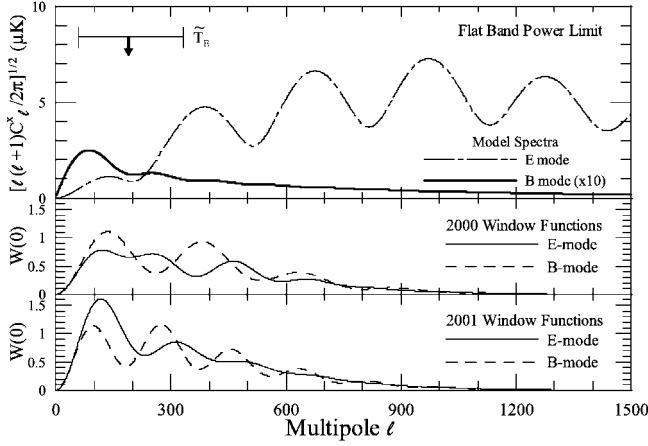
FIG. 2.—Zero-lag window functions for *E*- and *B*-modes (*solid and dashed lines*) for the 2000 observation season (*middle panel*) and for the 2001 season (*bottom panel*). The top panel shows the limit on T_E assuming $T_B \equiv 0$ for the combined *W*-band observations. For comparison, we also plot (*dashed line*) *E*-mode predictions from the best-fit model in Pryke et al. (2002) as well as *B*-mode predictions (shown in the top panel multiplied by 10) assuming the same model with $T/S = 1$.

Zaldarriaga (1998), the expression for $\langle \tilde{Q}_i \tilde{U}_j \rangle$ is given in terms of the *E*- and *B*-mode angular power spectra C_l^E and C_l^B by

$$\langle \tilde{Q}_i \tilde{U}_j \rangle = \sum_{lm} \frac{(2l+1)}{4\pi} [C_l^E + C_l^B] W_{12,l}(\phi_{ij}), \quad (3)$$

where $\phi_{ij} = \cos^{-1}(\hat{n}_i \cdot \hat{n}_j)$ is the lag. The window function $W_{12,l}$ has the form

$$W_{12,l} = \sum_m (B_{lm}^{QU})^2 F_{1,lm} F_{2,lm} \cos(m\phi_{ij}), \quad (4)$$

where the $F_{\{1,2\},lm}$ are given in terms of associated Legendre polynomials evaluated at the ring radius $\theta = 1°$. Here the beam function is

$$(B_{lm}^{QU})^2 = 4 \sin\left(\frac{m\pi}{2}\right) \cos\left(\frac{m\pi}{4}\right) \frac{\sin^2(m\pi/144)}{(m\pi/144)^2} e^{-l(l+1)\sigma^2}, \quad (5)$$

where $\sigma = 0°.10$ for the PIQUE beams. Similarly, the expression for $\langle \tilde{U}_i \tilde{U}_j \rangle$ is

$$\langle \tilde{U}_i \tilde{U}_j \rangle = \sum_l \frac{(2l+1)}{4\pi} [C_l^E W_{2,l}(\phi_{ij}) + C_l^B W_{1,l}(\phi_{ij})], \quad (6)$$

where $W_{1,l}$ and $W_{2,l}$ are the associated window functions given in H01. In this case, the beam function is given by

$$(B_{lm}^{UU})^2 = (B_{lm}^{QU})^2 \frac{\sin(m\pi/2)}{\cos(m\pi/4)}. \quad (7)$$

This formalism allows for combination of the two data sets; however, separate analyses of the $\tilde{Q}$ data and the $\tilde{U}$ data are also presented. We plot the zero-lag window functions in Figure 2.

The likelihood analysis proceeds by considering flat angular spectra, such that $l(l+1)C_l^X/2\pi = T_X^2$, where $X = E, B$. Since the amplitude C_l^B is predicted to be much smaller than C_l^E, we first find the limit $\tilde{T}_E$ on T_E under the assumption T_B is identically

TABLE 2
95% CONFIDENCE LIMITS

Year	Data[a]	$\tilde{T}_E$[b] (μK)	$\hat{T}_E$[c] (μK)	$\hat{T}_B$[c] (μK)
2000	CMB	12.7	15.8	14.7
	Quadrature	10.3	13.9	12.9
	(H1 − H2)/2	14.9	19.7	17.4
2001	CMB	10.4	15.9	17.8
	Quadrature	10.4	16.2	18.0
	(H1 − H2)/2	11.2	17.2	19.4
2000 + 2001	CMB	8.4	11.2	11.5
	Quadrature	6.1	8.3	8.4
	(H1 − H2)/2	8.5	11.5	11.6

[a] The two null tests are described in Table 1.
[b] The limit $\tilde{T}_E$ is found assuming $T_B \equiv 0$.
[c] The limits $\hat{T}_E$ and $\hat{T}_B$ are determined simultaneously by finding the contour of constant likelihood enclosing 95% of the volume.

zero.[5] This limit is compared to predictions in Figure 2. Next, joint upper limits ($\hat{T}_E$, $\hat{T}_B$) are determined by finding the constant contour of $L(T_E, T_B)$ enclosing 95% of the volume of L. This is repeated for each year separately and for the combination of the 2 years. The 95% CL upper limits are shown in Table 2, and the normalized likelihoods are shown in Figure 3. The null data sets are treated in an identical manner; results are tabulated in Table 2 and plotted in Figure 3. The limits in Table 2 do not include calibration errors. Note that for PIQUE's broad window functions, the 5% beam errors add only 2% errors in quadrature with the 10% calibration errors.

5. DISCUSSION

The main result here is a new constraint on the amount of polarized anisotropy in the CMB at subdegree angular scales. The result derives from combining data on *Q* from our first campaign with *U* from our second; in so doing, important information on the *Q-U* cross-correlation is included. We have summarized this result as a 95% CL limit of 8.4 μK on *E*-modes. Given PIQUE's window functions and current theoretical predictions (Fig. 3), we might expect a signal of a few microkelvins. The likelihood for our CMB data (Fig. 3, *bottom panel*) is consistent with this expectation. When we fit to an offset lognormal distribution (Bond, Jaffe, & Knox 2000), we

[5] The limit $\tilde{T}_E$ is found by integrating $L(T_E, 0)$; the result is ≈30% higher if $L(T_E^2, 0)$ is integrated.
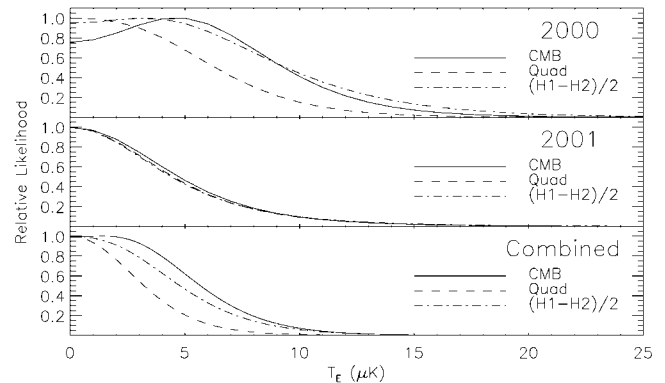

FIG. 3.—Normalized likelihoods vs. flat band power T_E under the assumption $T_B = 0$, as described in the text, for each year of data and for the combined data. In each panel, likelihoods for the CMB data are plotted along with likelihoods for two of the null data sets described in Table 1.

find a central value of 2.2 μK^2, a variance of $(16\ \mu K^2)^2$, and a noise-related offset of 18 μK^2.

The results presented here have been checked with two independent likelihood analyses and supported by extensive simulations. Our new selection criteria, as described above, are better able to deal with instrumental effects on different timescales, allow us to discard the 6 hr null test criterion we previously used, and work for both data sets together. However, applying the new criteria just to year 2000 data, we find a weaker 95% CL limit on polarized CMB anisotropy than in H01: 12.6 μK rather than 10.3 μK.

We have performed a variety of simulations to address the probability of such a change. Recall that our result for the year 2000 data was essentially unchanged by relaxing the null test cut and allowing in 80 extra hours of data, for a total of 330 hr. For these simulations, we start with roughly a 330 hr data set that is pure noise, generated assuming the actual weights for each period of the data set. We then investigate how cutting the data can change the derived limit. From this we find that although (1) the probability of the limit not worsening when 330 hr is reduced to 250 hr is only about 1%, (2) a 10 μK error from 330 hr of data is within 1 σ of what is expected from pure noise (given our experimental weights) and (3) the expected change in the limit from removing chunks of the data to get to 300 hr of data (our final sample for year 2000) is about 1.1 μK with a

standard deviation of about 1.6 μK. Thus, our observed change is again within 1 σ of what is expected.

We thus conclude that fluctuations alone can account for the change in the limit derived from the year 2000 data under different selection criteria. This is so even though the former cut on the null test, which we have shown is not needed, was probably too restrictive.

The expected signal, even including foregrounds, is still smaller than our new limit. Given the multipole range probed by PIQUE, this result provides the tightest constraint yet on the polarization spectrum predicted from primordial density fluctuations.

<div style="text-align:center">REFERENCES</div>

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